Constraints on the secular variation of Mercury’s magnetic field from the combined analysis of MESSENGER and Mariner 10 data

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Abstract Observations of Mercury’s internal magnetic field from the Magnetometer on the Mariner 10 spacecraft during its third flyby of Mercury (M10-III) in 1975 to constrain the secular variation in the internal field over the past 40 years. Although the second Mariner 10 flyby of Mercury was too distant from the planet to measure the internal field, the third Mariner 10 flyby (hereafter M10-III) on 16 March 1975 confirmed the existence of an intrinsic magnetic field. The M10-III flyby had a closest approach distance of 327 km at geographic latitude 68°N, and analyses of the magnetic field data yielded a dipole moment of (136–350) nT Rm3, where Rm is Mercury’s radius (2440 km) and where the range of estimates reflected uncertainties in the external fields and the nondipole components of the internal field [Ness et al., 1975; Connerney and Ness, 1988]. Several hypotheses for the origin of Mercury’s field were suggested, including a core dynamo field [e.g., Connerney and Ness, 1988] and remanent crustal fields [e.g., Stephenson, 1976; Aharonson et al., 2004].

1. Introduction

Magnetic field data acquired by the Mariner 10 spacecraft during its first flyby of Mercury in March 1974 [Ness et al., 1974] suggested that Mercury has a terrestrial-like magnetosphere with a weak intrinsic magnetic field. However, a definitive assessment of the geometry and origin of the internal field was not possible because of the equatorial spacecraft trajectory, the large distance of closest approach, and the substantial field variability on the outbound portion of the flyby [Ness et al., 1974]. Although the second Mariner 10 flyby of Mercury was too distant from the planet to measure the internal field, the third Mariner 10 flyby (hereafter M10-III) on 16 March 1975 confirmed the existence of an intrinsic magnetic field. The M10-III flyby had a closest approach distance of 327 km at geographic latitude 68°N, and analyses of the magnetic field data yielded a dipole moment of (136–350) nT Rm3, where Rm is Mercury’s radius (2440 km) and where the range of estimates reflected uncertainties in the external fields and the nondipole components of the internal field [Ness et al., 1975; Connerney and Ness, 1988]. Several hypotheses for the origin of Mercury’s field were suggested, including a core dynamo field [e.g., Connerney and Ness, 1988] and remanent crustal fields [e.g., Stephenson, 1976; Aharonson et al., 2004]. More than three decades later, the Magnetometer on the MESSENGER spacecraft has provided measurements of Mercury’s magnetic field environment during flybys in 2008 and 2009 and from orbit around the planet since March 2011. The flyby data confirmed the dominantly dipolar geometry of the internal field and indicated a core dynamo origin for the field [Anderson et al., 2008; Purucker et al., 2009; Uno et al., 2009]. Orbital data have enabled the development of a time-averaged model for Mercury’s magnetic field. The model includes an average dipole moment of 190 nT Rm3 offset 0.2 Rm northward from the planetary equator [Anderson et al., 2011, 2012; Johnson et al., 2012] and aligned with the planetary rotation axis. Mercury’s weak dipole moment, the northward offset of the magnetic equator from the planetary equator, and the axisymmetry of the field about the rotation axis yield a spherical harmonic description for the internal field dominated by the axial dipole term, g1,0, but with important contributions from zonal harmonics of the next three highest degrees, g2,0, g3,0, and g4,0 [Anderson et al., 2012]. This last result arises because the
spherical harmonic expansion for an offset dipole includes not only dipole and quadrupole ($g_2^0$) terms [Bartels, 1936] but higher-degree terms as well. The low-degree zonal terms provide first-order constraints on dynamo models for field generation in Mercury. Models proposed for weak dipole-dominated internal fields invoke aspects of the geometry of the fluid outer core, the presence of one or more stably stratified layers in the outer core, or feedback between the magnetospheric and core fields [Stanley and Glatzmaier, 2010]. Current dynamo modeling efforts are focused on capturing the asymmetry in the field about the geographic equator [e.g., Cao et al., 2014].

The separation of nearly 40 years between the Mariner 10 and MESSENGER observations provides an opportunity to examine temporal variability (secular variation) in Mercury’s internal magnetic field. Secular variation, and in particular the spatial power spectrum of secular variation, provides an additional important constraint on core dynamo models. For example, models that invoke a stably stratified layer at the top of the core to produce a weak field at and above the planetary surface also predict suppressed secular variation relative to models with no such layer [e.g., Christensen, 2006]. In this paper, we reanalyze the Mariner 10 and MESSENGER flyby data using techniques developed for the analysis of MESSENGER orbital data and search for evidence of secular variation in the planetary field. Given the limitations of the magnetic field data from the first Mariner 10 flyby, we consider data only from the third flyby. We first identify the instantaneous position of the magnetic equator in the M10-III data and compare this position to the range of values seen in the MESSENGER orbital data. We then use the magnetospheric model of Johnson et al. [2012] to examine any changes in the dipole moment or dipole offset. Finally, we briefly revisit the first two MESSENGER flybys. Taken together, the Mariner 10 and MESSENGER mission data sets constrain estimates of secular variation in the axial dipole, quadrupole, and octupole components of Mercury’s dynamo field.


On its third flyby of Mercury, Mariner 10 entered the magnetosphere in the near-tail region and passed over the planet with a closest approach at an altitude of approximately 0.14 $R_M$ before exiting the magnetosphere near dawn (Figure 1). We used 1.2 s averages of the magnetic field measurements (together with the standard deviation evaluated over 1.2 s) taken from Lepping et al. [1979] and the Planetary Data System to identify inner and outer limits for the bow shock and magnetopause boundary crossings following the method used for MESSENGER orbital data [Winslow et al., 2013]. The resulting crossing times agree well with values given by Ness et al. [1976] and Lepping et al. [1979] and are shown in Figure 1.
The northward offset of Mercury’s internal dipole field has been identified by the position of the magnetic equator in the MESSENGER orbital magnetic field data [Anderson et al., 2011, 2012]. The geometry of MESSENGER’s orbit around Mercury results in covariance among terms of even and odd degree in a spherical harmonic expansion for the internal field, similar to the covariance structure for M10-III [Connerney and Ness, 1988; Korth et al., 2004]. Identification of the magnetic equator resolves this fundamental trade-off between the dipole and quadrupole terms. In Mercury solar orbital (MSO) coordinates (in which \(X_{\text{MSO}}\) is sunward, \(Z_{\text{MSO}}\) is northward and aligned with the planetary spin axis, and \(Y_{\text{MSO}}\) completes the right-handed system), the instantaneous location of the magnetic equator is given by the \(Z_{\text{MSO}}\) position at which \(B_\rho = 0\), where \(B_\rho\) is the radial component of the magnetic field in MSO cylindrical coordinates. MESSENGER crossings of the magnetic equator between 24 March 2011 and 2 March 2012 yielded an average northward offset from the geographic equator of 479 ± 6 km for descending (low-altitude) crossings and 486 ± 74 km for ascending (high-altitude) crossings [Anderson et al., 2012].

Following the same procedure, we identified the equator crossing region in the M10-III data. The spacecraft \(Z_{\text{MSO}}\) position as a function of \(B_\rho\) was fit by a quadratic function, and the instantaneous magnetic equator position, \(Z_{\rho_0}\), was estimated by the intercept of the fit \(B_\rho = 0\) (Figure 2a). The M10-III data yielded an equator offset of \(Z_{\rho_0} = 882 ± 35\) km, where the uncertainty estimate is three standard errors in the intercept of the quadratic fit.

The instantaneous equator position in the M10-III data is significantly northward of the average calculated from the MESSENGER orbital data. The M10-III equator crossing occurs at an altitude of 3415 km and an MSO longitude of 113°E, comparable to some of the high-altitude MESSENGER crossings. For comparison, we analyzed magnetic equator crossings identified at high altitudes on ascending legs of MESSENGER orbits between 24 March 2011 and 15 December 2012 at MSO longitudes ranging from 80°E to 125°E (Figure 2b). These crossings lie between 746 km south of the planetary equator and 1058 km north of the planetary equator. The M10-III equator offset falls within this range and is thus not atypical of instantaneous crossings from MESSENGER orbits with trajectories most similar to that of the M10-III flyby trajectory. The large variation in the MESSENGER crossings, particularly in the far tail, has been suggested to reflect tilting of the tail in response to north-south excursions in the solar wind velocity [Anderson et al., 2012]. In addition, the M10-III equator crossing occurred within 1 min of the inbound inner magnetopause crossing and may therefore have been distorted by the near field of magnetopause currents.

### 3. Estimation of Dipole Moment

The time-averaged model for Mercury’s magnetosphere developed by Johnson et al. [2012] was derived from MESSENGER orbital data between 24 March 2011 and 12 December 2011 and used the paraboloid...
The magnetopause standoff distance, \( R_{SS} \), can vary substantially over time scales of minutes [Winslow et al., 2013]. The average outbound (dayside) magnetopause crossing yields an \( R_{SS} \) of 1.21 \( R_M \). A similar calculation using the average inbound (nightside) crossing gives an \( R_{SS} \) of 1.40 \( R_M \). This difference in \( R_{SS} \) values for the inbound and outbound crossings is not atypical of MESSENGER orbital observations. On time scales of
Mercury’s orbital period, \( R_{SS} \), also varies with Mercury’s heliocentric distance, \( r_{helio} \) [Korth et al., 2012; Johnson et al., 2013]. MESSENGER orbital observations indicate an average relation of the form \( R_{SS} = 1.99 \left( r_{helio} \right)^{1/3} \). This relation yields an \( R_{SS} \) value of 1.54 \( R_M \) given a heliocentric distance of 0.356 AU at the time of the M10-III flyby.

We assessed the impact of our choice of \( R_{SS} \) on our estimate of dipole moment as follows: for each of the \( R_{SS} \) values of 1.40 \( R_M \) and 1.54 \( R_M \), we held the dipole offset constant at the best fit value of 475 km obtained above, and we found a best fit dipole moment by minimizing the misfit between the magnetospheric model

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**Figure 3.** (a) RMS misfit between model and M10-III data obtained from a grid search over dipole moment and offset of the dipole from the geographic equator. Contours indicate misfits from 10 nT to 60 nT in 10 nT steps. The parameter pairs for the best fit offset of \( Z_d = 475 \) km and the M10-III magnetic equator offset of 882 km are marked by the filled black circle and filled square, respectively. The diagonal dashed line shows the trade-off between best fit dipole offset and dipole moment; blue triangles correspond to parameter pairs that yield an RMS misfit 25% above the minimum value. The horizontal dashed line shows the location of an RMS misfit profile corresponding to the best fit dipole offset; a 25% increase above the minimum misfit would yield dipole moments shown by the black triangles. (b) Magnetic field observations from M10-III inside the magnetopause (solid lines) and best fit models to the data for dipole offsets of 475 km (dashed line) and 882 km (dotted line).
and the M10-III data. The resulting best fit dipole moment was insensitive to the choice of $R_{SS}$, varying by only 2 nT $R_M^3$ from the value obtained using the initial $R_{SS}$ estimate of 1.21 $R_M$.

### 3.2. MESSENGER Flybys

We reanalyzed the first two MESSENGER flybys of Mercury (hereafter M1 and M2, respectively) on 14 January and 6 October 2008 to investigate whether the magnetic field measured during those flybys is consistent with the time-averaged magnetospheric model determined from MESSENGER orbital data [Johnson et al., 2012]. The negligible latitudinal variations during the M1 and M2 equatorial trajectories meant that it was not possible to identify a magnetic equator crossing from data acquired during either flyby.

We set all parameters in the model to the time-averaged values except the magnetopause standoff distance, $R_{SS}$, and the dipole moment, $m$. Magnetopause crossing positions were taken from Anderson et al. [2008] and Slavin et al. [2009] for M1 and M2, respectively. For both flybys the $R_{SS}$ estimates from the $R_{SS} - \mu_{helio}$ relation were found to be very similar to the $R_{SS}$ value calculated from the outbound (dayside) magnetopause crossing position. The corresponding grid searches over dipole moment also gave similar results, so we consider the outbound $R_{SS}$ values and the associated dipole moment searches to best represent the flyby data (Table 1). The best fit dipole moment of 188 nT $R_M^3$ for the combined M1 and M2 data is consistent with the time-averaged magnetospheric model of Johnson et al. [2012].

### 4. Discussion and Conclusions

Magnetic field observations from the third Mariner 10 flyby and the MESSENGER flybys of Mercury have been analyzed using approaches applied to the MESSENGER orbital data to investigate changes in the dipole moment and offset over the intervening four decades. The magnetic field measured along the M10-III trajectory was dominated by contributions from the magnetopause and the internal dipole field. Paraboloid magnetospheric models were generated in which the magnetopause fields were prescribed using $R_{SS} = 1.21 R_M$ and a grid search was conducted over dipole moment, $m$, and dipole offset from the geographic equator, $Z_d$. The best fit model is indistinguishable from that of the MESSENGER orbital data (Table 1). The best fit value for the dipole offset yields a substantially lower misfit than that obtained using an offset given by the magnetic equator crossing indicated by the M10-III data. This result supports the conclusion that the instantaneous value is unlikely to be representative of the average dipole offset. The magnetic equator position indicated by the M10-III data, although farther north than the MESSENGER average, is well within the observed range in MESSENGER orbital data and likely an indication of a tilted current sheet or perhaps substantial magnetopause fields rather than a high offset of instantaneous magnetic equator. The first two MESSENGER flybys also yield best fit moments consistent with that obtained from MESSENGER orbital data.

Although the best fit models yield no resolvable evidence for secular variation in either the dipole moment or its offset, the trade-off between dipole moment and offset provides upper bounds on secular variation in the axisymmetric part of the field. We find that a 25% increase in RMS misfit results in a noticeably worse fit to the data, and we take the $(m, Z_d)$ values of (206 nT $R_M^3$, 370 km) and (173 nT $R_M^3$, 560 km) corresponding to such a misfit as upper limits on secular variation. The equivalent spherical harmonic representation of this offset dipole field can be found with the relationships given in Table 3 of Anderson et al. [2012]. The above $(m, Z_d)$ limits correspond to low-degree axial spherical harmonic coefficients $g_1^0$, $g_2^0$, and $g_3^0$ of (−206, −62, −14) nT and (−173, −79, −27) nT, respectively, and the mean values from the MESSENGER orbital data are (−190, −75, −22) nT [Anderson et al., 2012]. Thus, changes in the dipole moment (or the $g_1^0$ coefficient) of less than 10% may be accompanied by changes in the degree 2 and degree 3 axial terms of up to 16% and 35%, respectively, yielding substantial changes in the slope of the low-degree spatial power spectrum. We did not consider the $g_3^3$ term here, as its value indicated by MESSENGER orbital data is small, $|g_3^3| < 6$ nT.

The results reported here were derived under the assumption that there is no contribution from nonaxisymmetric low-degree terms in the field. We verified that there is no substantial tilt of the dipole axis from the spin axis as follows. We subtracted the magnetopause and tail fields of the best fit model from the M10-III data sampled at 6 s intervals. On the assumption that the remaining signal is primarily that of an offset dipole, we calculated the apparent latitude and longitude of the pole from the magnetic field direction at each point along the trajectory. The pole latitude is measured on a sphere centered on the offset dipole ($Z_d = 475$ km). Pole positions at high altitudes were clearly affected by errors in the model estimate of the
magnetopause fields. For data acquired at altitudes less than 650 km, the pole positions cluster about a mean position of 87.9°N, 78.8° E. A magnetospheric model fit with a dipole tilt of 2.1° in the azimuth direction 78.8°E gave less than a 3% improvement (0.3 nT) in RMS misfit over the zero-tilt model. The 95% confidence limit on the pole position is 0.7° if all 42 estimates of the pole position at 6 s sampling are considered to be independent. However, larger confidence limits result if all of the observations are not taken to be independent. MESSENGER orbital data indicate an upper bound on the current dipole tilt of 0.8° via analysis of the magnetic equator crossings, and constraints on the tilt direction are poor [Anderson et al., 2012]. We conclude that there is no evidence of significant secular variation in Mercury’s dipole tilt. Furthermore, we verified via fits to the paraboloid magnetospheric model that any dipole offset in the plane of the geographic equator is insignificant in comparison with the offset in the Z direction. Taken together, these results confirm that there has been no observable secular variation in the nonaxial contributions to the field at low spherical harmonic degree and order.

The presence of secular variation in Earth’s magnetic field is well documented [e.g., Finlay et al., 2010]. Studies of the magnetic fields of Jupiter and Saturn have revealed no definitive secular variation in the internal field of either body over more than 30 years of observations [Yu et al., 2010; Cao et al., 2011; Ridley, 2012]. It is interesting to consider whether Earth-like secular variation would be detectable in the Mercury magnetic field data. Between 1975 and 2010, secular variation in the terrestrial axial dipole, $g_1^0$, axial quadrupole, $g_2^0$, and nonaxial dipole terms, $g_1^1$ and $h_1^1$, resulted in changes in these terms of 2%, 21%, 27%, and 15%, respectively [Finlay et al., 2010]. At Mercury, a 2% change in the dipole moment corresponds to a variation of less than 4 nT $R_M^3$, and Earth-like percentage variations in $g_1^1$ and $h_1^1$ would be difficult to assess via changes in the pole position, given the negligible tilt of the dipole moment from the spin axis. From our upper limits on changes in tilt and the low-degree axial terms at Mercury, Earth-like relative magnitudes of secular variation are permitted but not required by the M10-III data. From a dynamical standpoint, the spatial power spectrum of secular variation at Mercury may be quite different from that at Earth. Dynamo models that match the weak, equatorially asymmetric field with a stably stratified layer at the top of the core or through a combination of volumetric buoyancy and equatorially symmetric core-mantle heat flow have characteristic timescales for secular variation in the low-degree terms that are of order 1000 years, and the amplitude of the variations is small [Christensen and Wicht, 2008; Cao et al., 2014]. Strong secular variation in $g_1^0$ and $g_2^0$ (±100% of the peak values) occurs in oscillatory dynamos (e.g., Dietrich and Wicht [2013], though these models were not specifically adapted to Mercury), but the oscillatory time scale of a few thousand years results in very little change in these coefficients over the 40 year interval between the times of the third Mariner 10 flyby and MESSENGER orbital observations. Our results provide the first observational constraint on secular variation in Mercury’s internal magnetic field at low spherical harmonic degree.

References


